

ELEVEN EXOPLANET HOST STAR ANGULAR DIAMETERS FROM THE CHARA ARRAY

ELLYN K. BAINES^{1,4}, HAROLD A. MCALISTER², THEO A. TEN BRUMMELAAR², JUDIT STURMANN², LASZLO STURMANN²,
NILS H. TURNER², AND STEPHEN T. RIDGWAY³

¹ Remote Sensing Division, Naval Research Laboratory, 4555 Overlook Avenue SW, Washington, DC 20375, USA; ellyn.baines.ctr@nrl.navy.mil

² Center for High Angular Resolution Astronomy, Georgia State University, P.O. Box 3969, Atlanta, GA 30302-3969, USA; hal@chara.gsu.edu;
theo@chara-array.org, judit@chara-array.org, sturmman@chara-array.org, nils@chara-array.org

³ Kitt Peak National Observatory, National Optical Astronomy Observatory, P.O. Box 26732, Tucson, AZ 85726-6732, USA; ridgway@noao.edu

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ABSTRACT

We directly measured the angular diameters for 11 exoplanet host stars using Georgia State University’s CHARA Array interferometer and calculated their linear radii and effective temperatures. The sample tends toward evolving or evolved stars and includes one dwarf, four subgiants, and six giants. We then estimated masses and ages for the stars using our effective temperatures combined with metallicity measurements from the literature.

Key words: infrared: stars – planetary systems – stars: fundamental parameters – techniques: interferometric

1. INTRODUCTION

Exoplanets are discovered on a regular basis, most via radial velocity surveys and transiting events. Many host star angular diameters have been estimated using photometric and spectroscopic methods (e.g., Ribas et al. 2003; Fischer & Valenti 2005, respectively), and while these are excellent for approximating angular diameters, they are by nature indirect methods. The advantage interferometry brings is the ability to directly measure the angular sizes of the stars, which in turn leads to physical radii and effective temperatures. These are important parameters that describe the parent star as well as the environment in which the exoplanet resides.

This paper represents an extension and continuation of the work described in Baines et al. (2008), where the angular diameters for 24 exoplanet host stars were published. While the previous sample featured a few giants and some subgiants, well over half were dwarfs or stars showing signs of just beginning to evolve off the main sequence. This paper focuses on giants and subgiants, and only one dwarf is represented.

2. INTERFEROMETRIC OBSERVATIONS

All observations were obtained using the Center for High Angular Resolution Astronomy (CHARA) Array, a six-element optical/infrared interferometric array located on Mount Wilson, California (ten Brummelaar et al. 2005). We used the pupil-plane “CHARA Classic” beam combiner in the K' -band ($2.15 \mu\text{m}$), paired with the longest baseline the Array offers at 331 m. The observing procedure and data reduction process employed here are described in McAlister et al. (2005). Table 1 lists the exoplanet host stars observed, their calibrators, the dates of the observations, and the number of observations obtained.

Our target list was culled from the complete exoplanet list by using declination limits and magnitude constraints: north of -10° declination, brighter than $V = +10$ in order for the tip/tilt system to lock onto the star, and brighter than $K = +6.5$, so fringes were easily visible. We obtained data on the 11 exoplanet host stars over two observing runs in July and September 2008.

Reliable calibrator stars are critical in interferometric observations, acting as the standard against which the science target is measured, so every effort was made to find spherical, non-variable, single-star calibrators. Our observing pattern was calibrator-target-calibrator so that every target was flanked by calibrator observations made as close in time as possible; therefore “10 bracketed observations” denotes 10 object and 11 calibrator data sets, each of which is comprised of approximately 200 scans across the fringe. This allowed us to calculate the target’s calibrated visibilities from the instrumental visibilities of the target and calibrator. Figure 1 shows an example of uncalibrated visibilities. Acceptable calibrators were chosen to be smaller than ~ 0.4 milliarcseconds (mas), so they were nearly unresolved and uncertainties in their diameters did not affect the target’s diameter calculation as much as if the calibrator had a significant angular size on the sky.

In order to estimate the calibrator stars’ angular diameters as well as check for excess emission that could indicate a low-mass stellar companion or circumstellar disk, we fitted spectral energy distributions (SEDs) based on published *UBVRIJHK* photometric values for each star. Limb-darkened (LD) diameters were calculated using Kurucz model atmospheres⁵ based on effective temperature (T_{eff}) and gravity ($\log g$) values obtained from the literature. The models were then fit to observed photometric values also from the literature after converting magnitudes to fluxes using Colina et al. (1996) for *UBVRI* values and Cohen et al. (2003) for *JHK* values.

Table 1 lists the T_{eff} and $\log g$ used for each calibrator, the resulting LD angular diameters, and the distance between the target and calibrator stars. We used calibrators as close to the target star as possible. The target-calibrator (T-C) distances ranged from 1° to 9° and all but two calibrators were within 5° of their target stars. This allowed us to observe the stars as close together in time as possible, usually on the order of 3 to 5 minutes between the two. For the T-C pairs of 8° and 9° , the slightly greater distance added little to the error in the diameter measurement. Table 2 provides more details on each calibrator star used, and Table 3 lists the Modified Julian Date (MJD), projected baseline (B), projected baseline position angle (Θ), calibrated visibility (V_c), and error in V_c (σV_c) for each exoplanet host star observed.

⁴ The observations described here were completed while with the Center for High Angular Resolution Astronomy, Georgia State University, P.O. Box 3969, Atlanta, GA 30302-3969, USA.

⁵ See <http://kurucz.cfa.harvard.edu>.

Table 1
Observing Log and Calibrator Stars' Basic Parameters

Observing Log					Calibrator Information			
Target HD	Other Name	Calibrator HD	Date (UT)	No. of Bracketed Observations	T_{eff} (K)	$\log g$ (cm s^{-2})	$\theta_{\text{LD,SED}}$ (mas)	T-C Sep (deg)
16141	79 Cet	18331	2008 Sep 9	10	8710	4.14	0.354 ± 0.019	5
17092	...	14212	2008 Sep 11	5	9333	4.08	0.291 ± 0.006	5
45410	6 Lyn	46590	2008 Sep 11	5	9550	4.14	0.221 ± 0.007	2
154345	...	151044	2008 Sep 10	7	6166	4.38	0.380 ± 0.008	4
185269	...	184381	2008 Jul 18	15	6650	4.34	0.285 ± 0.010	3
			2008 Jul 20	5
188310	ξ Aql	182101	2008 Sep 8	8	6607	4.33	0.344 ± 0.014	8
199665	18 Del	194012	2008 Sep 8	10	6310	4.36	0.441 ± 0.016	9
210702	...	210074	2008 Sep 8	4	7079	3.82	0.384 ± 0.013	4
217107	...	217131	2008 Sep 8	5	6918	3.71	0.305 ± 0.014	1
221345	14 And	222451	2008 Sep 11	5	6761	4.22	0.346 ± 0.011	3
222404	γ Cep	219485	2008 Jul 17	3	9790	4.14	0.214 ± 0.006	4
			2008 Sep 11	7

Notes. T_{eff} and $\log g$ values come from Allende Prieto & Lambert (1999), except for HD 184381 and HD 219485, whose T_{eff} and $\log g$ values are based on spectral type as listed in the *SIMBAD Astronomical Database* and Cox (2000).

Table 2
Previous Calibrator Uses

Calib HD	Previous Calibrator Uses
14212	Used as calibrator in van Belle & von Braun (2009)
18331	Used as calibrator in van Belle & von Braun (2009)
46590	Considered a single star in Royer et al. (2007)
151044	Used as calibrator in Baines et al. (2008)
182101	Used as calibrator in Berger et al. (2006)
184381	Used as calibrator in Johnson et al. (2006)
194012	Used as calibrator in Baines et al. (2008) & Montes et al. (1995); no binary companion found in McAlister et al. (1987)
210074	Used as comparison star in Wittenmyer et al. (2005) & Henry et al. (2000)
217131	Used as comparison star in Vogt et al. (2005); no binary companion found in McAlister et al. (1987)
219485	Considered a single star in Royer et al. (2007)
222451	Considered a single star in Nordström et al. (2004)

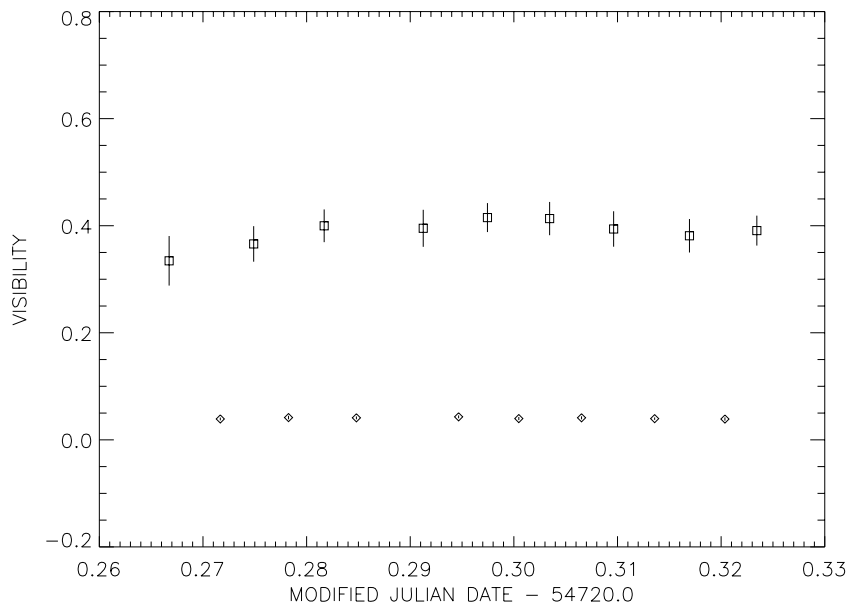


Figure 1. Uncalibrated visibilities for HD 222404 from 2008 September 11. The squares and diamonds are the calibrator's and target's measured visibilities, respectively, and the vertical lines are the errors in those visibilities.

Table 3
Calibrated Visibilities

Target HD	MJD	B (m)	Θ (deg)	V_c	σV_c	
16141	54718.438	285.47	237.1	0.915	0.134	
	54718.445	281.55	238.2	0.944	0.123	
	54718.450	278.72	239.0	0.845	0.113	
	54718.455	275.69	240.0	0.961	0.100	
	54718.461	272.09	241.2	0.815	0.104	
	54718.467	269.11	242.2	0.834	0.115	
	54718.472	266.31	243.3	0.900	0.119	
	54718.477	263.67	244.4	0.972	0.143	
	54718.482	260.84	245.6	0.845	0.146	
	54718.487	258.06	246.8	0.977	0.138	
	17092	54720.344	285.50	221.5	0.904	0.107
		54720.354	290.19	223.3	0.841	0.093
		54720.364	295.06	225.3	0.811	0.081
		54720.371	297.63	226.5	0.753	0.079
		54720.380	301.41	228.4	0.846	0.075
45410	54720.481	258.11	212.9	0.696	0.078	
	54720.490	263.24	215.0	0.651	0.053	
	54720.496	266.69	216.4	0.587	0.073	
	54720.502	269.68	217.7	0.665	0.106	
	54720.509	272.90	219.2	0.716	0.097	
154345	54719.168	328.79	90.5	0.885	0.094	
	54719.179	328.73	93.3	0.843	0.109	
	54719.185	328.66	94.7	0.811	0.089	
	54719.192	328.57	96.2	0.803	0.096	
	54719.198	328.45	97.6	0.847	0.096	
	54719.204	328.29	99.2	0.903	0.095	
	54719.213	328.00	101.4	0.817	0.122	
	185269	54665.204	321.00	228.6	0.860	0.146
54665.216		323.97	230.0	0.946	0.129	
54665.226		326.17	231.3	0.757	0.148	
54665.236		327.81	232.6	0.926	0.110	
54665.245		328.96	233.9	0.928	0.178	
54665.404		323.06	266.1	0.771	0.064	
54665.410		322.92	267.7	0.741	0.050	
54665.417		322.85	269.2	0.816	0.048	
54665.423		322.85	90.8	0.921	0.057	
54665.430		322.93	92.4	0.877	0.075	
54665.438		323.11	94.3	0.912	0.084	
54665.445		323.35	96.0	0.910	0.091	
54665.452		323.68	97.7	0.855	0.080	
54665.459		324.06	99.4	0.927	0.083	
54665.466		324.52	101.1	0.841	0.129	
54667.381		323.73	262.0	1.004	0.096	
54667.387		323.44	263.5	0.830	0.103	
54667.393		323.21	264.9	0.892	0.096	
54667.400		323.02	266.5	1.014	0.085	
54667.406		322.90	267.9	0.899	0.113	
188310	54717.211	293.54	249.5	0.103	0.014	
	54717.223	289.87	252.2	0.106	0.017	
	54717.229	288.10	253.7	0.107	0.012	
	54717.236	286.11	255.5	0.106	0.014	
	54717.242	284.72	257.0	0.094	0.015	
	54717.248	283.37	258.5	0.110	0.019	
	54717.253	282.29	260.0	0.111	0.018	
	54717.259	281.25	261.6	0.127	0.018	
199665	54717.336	285.96	90.6	0.614	0.064	
	54717.341	286.09	92.1	0.567	0.062	
	54717.347	286.36	93.6	0.562	0.077	
	54717.352	286.78	95.0	0.574	0.053	
	54717.358	287.37	96.6	0.566	0.060	
	54717.364	288.15	98.2	0.512	0.055	
	54717.370	289.11	99.1	0.479	0.069	
	54717.377	290.31	101.5	0.482	0.049	
	54717.383	291.58	103.1	0.414	0.035	
	54717.390	293.30	104.9	0.500	0.065	

Table 3
(Continued)

Target HD	MJD	B (m)	Θ (deg)	V_c	σV_c
210702	54717.426	302.96	100.6	0.635	0.076
	54717.436	304.66	103.1	0.652	0.072
	54717.442	305.68	104.5	0.591	0.085
217107	54717.448	306.87	105.9	0.640	0.091
	54717.283	292.41	236.3	0.771	0.096
	54717.289	289.09	237.2	0.793	0.127
	54717.296	285.35	238.3	0.757	0.095
	54717.303	281.40	239.5	0.799	0.118
221345	54717.309	278.11	240.6	0.776	0.114
	54720.234	313.74	229.1	0.278	0.031
	54720.239	315.41	229.9	0.253	0.034
	54720.245	317.13	230.8	0.266	0.028
	54720.250	318.64	231.7	0.232	0.024
222404	54720.256	320.12	232.7	0.251	0.028
	54664.457	253.07	230.4	0.105	0.011
	54664.466	254.63	233.0	0.099	0.011
	54664.475	256.07	235.6	0.091	0.010
	54720.278	247.87	222.5	0.104	0.012
	54720.285	249.26	224.5	0.093	0.010
	54720.295	251.32	227.6	0.093	0.008
	54720.301	252.45	229.3	0.086	0.008
	54720.307	253.58	231.2	0.092	0.009
	54720.313	254.70	233.2	0.091	0.008
54720.320	255.83	235.2	0.087	0.009	

Notes. The projected baseline position angle (Θ) is calculated to be east of north.

3. ANGULAR DIAMETER DETERMINATIONS

Diameter fits to visibilities (V) were based upon the uniform disk (UD) approximation given by $V = [2J_1(x)]/x$, where J_1 is the first-order Bessel function and $x = \pi B \theta_{UD} \lambda^{-1}$, where B is the projected baseline at the star's position, θ_{UD} is the apparent UD angular diameter of the star, and λ is the effective wavelength of the observation (Shao & Colavita 1992). The LD relationship incorporating the linear limb darkening coefficient μ_λ (Hanbury-Brown et al. 1974) is

$$V = \left(\frac{1 - \mu_\lambda}{2} + \frac{\mu_\lambda}{3} \right)^{-1} \times \left[(1 - \mu_\lambda) \frac{J_1(x)}{x} + \mu_\lambda \left(\frac{\pi}{2} \right)^{1/2} \frac{J_{3/2}(x)}{x^{3/2}} \right]. \quad (1)$$

Figures 2 and 3 show the LD diameter fits for all the stars. Though the difference between LD and UD diameters is a minor effect in the wavelength used here, the former have the advantage over the latter in that they are better suited to calculating effective temperatures and more closely represent the physical properties of the star (van Belle & von Braun 2009).

For each θ_{LD} fit, the errors were derived via the reduced χ^2 minimization method: the diameter fit with the lowest χ^2 was found and the corresponding diameter provided the final θ_{LD} for the star. The errors were calculated by finding the diameter at $\chi^2 + 1$ on either side of the χ^2_{min} and determining the difference between the χ^2 diameter and $\chi^2 + 1$ diameter.

Our experience has shown that the root-mean-square (rms) of the residuals to diameter fits of visibilities is typically smaller than the mean of the standard errors attributed to each contributing visibility measurement. As described by McAlister et al. (2005), the error estimates assigned to calibrated visibilities were determined by the rms of the means of subsets of the entire sample of visibility measurements made at a particular epoch.

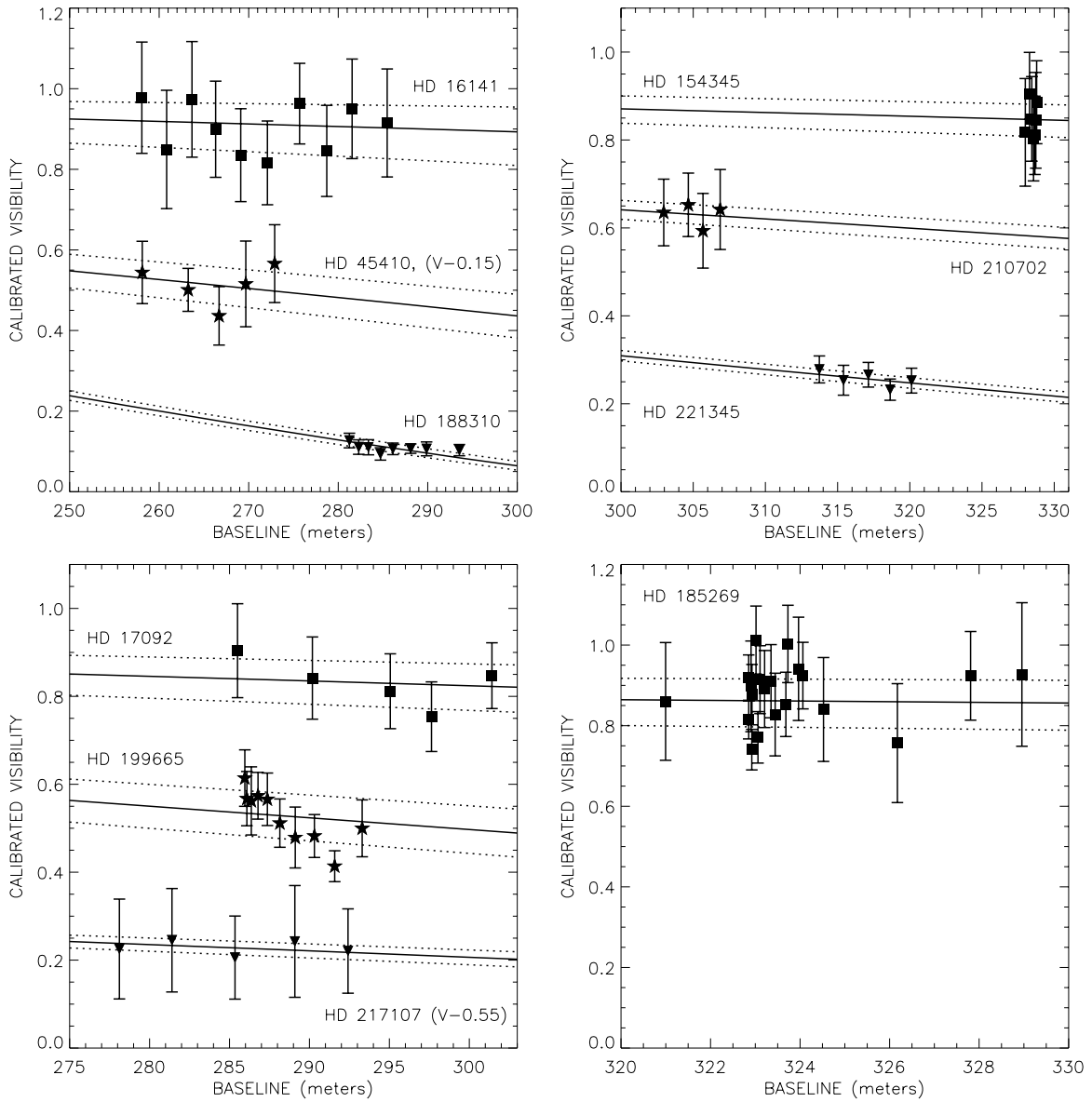


Figure 2. LD disk diameter fits for all the stars except HD 222404. The solid lines represent the theoretical visibility curve with the best fit θ_{LD} for each star, the dashed lines are the 1σ error limits of the diameter fit, the solid symbols are the calibrated visibilities, and the vertical lines are the measured errors. HD 45410's and HD 217107's visibilities were subtracted by the offset indicated by "(V - #)" so they would not overlap other data points.

We now find that this approach tends to overestimate the error of individual visibilities, producing reduced χ^2 values well under 1.0. This, in turn, leads to overestimates of the errors in angular diameter. In calculating the diameter errors in Table 4, we have adjusted the estimated visibility errors by a factor that forces the reduced χ^2 to unity, and we believe the resulting diameter errors are more representative of the influence of the true intrinsic errors in our visibilities.

Table 4 lists the following parameters for each star: spectral type, μ_λ , the *Hipparcos* parallax (π , van Leeuwen 2007), the LD diameter estimated from SED fits (θ_{SED}), the UD and LD angular diameters θ_{UD} and θ_{LD} , and the linear radius (R_L) derived from the combination of θ_{LD} and π . Six of the stars had θ_{SED} calculated by van Belle & von Braun (2009), and Table 4 lists the photometric sources for the remaining stars, whose SED fits were completed by us as described in Section 2. The T_{eff} and $\log g$ values used in our SED fits were from Allende Prieto & Lambert (1999) for all the stars except HD 17092 and HD

154345, which were from Cox (2000) and Valenti & Fischer (2005), respectively. The star HD 17092 does not have any available parallax measurements, so we used the photometric distance estimate from Gontcharov (2008) with an assigned error of 10%.

To check how well the estimated angular diameters match the measured values, Figure 4 plots θ_{SED} versus θ_{LD} and shows how the SED diameters slightly underestimate the true sizes of these evolved stars. This may be due to model assumptions about opacity that are not exactly true to life.

Two stars have been previously measured interferometrically: HD 221345 and HD 222404. van Belle et al. (1999) observed HD 221345 using the Palomar Testbed Interferometer (Colavita et al. 1999) and their value of θ_{UD} was 1.75 ± 0.07 mas. Nordgren et al. (1999) used the Navy Prototype Optical Interferometer (NPOI; Armstrong et al. 1998) to measure HD 222404 and their LD angular diameter of 3.24 ± 0.03 mas is close to our measurement of 3.30 ± 0.01 mas. The NPOI observes in visible

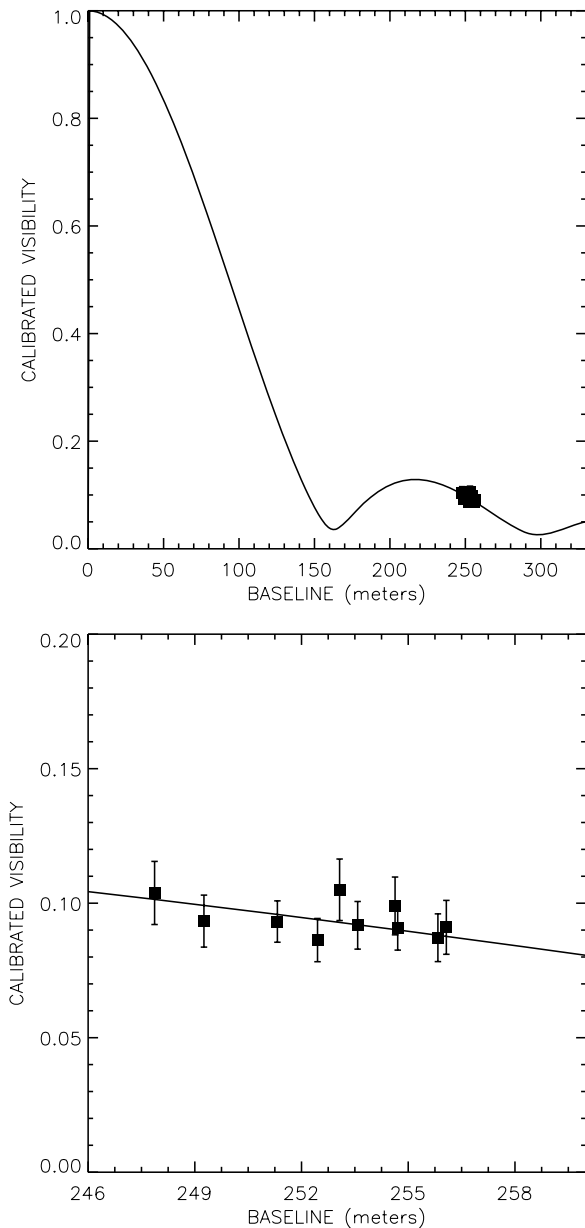


Figure 3. LD disk diameter fit for HD 222404. The solid line represents the theoretical visibility curve for the best fit θ_{LD} , the squares are the calibrated visibilities, and the vertical lines are the measured errors. The top panel shows the full visibility curve with the 10 data points clustered on the second lobe, and the bottom panel zooms in on those data points.

wavelengths and therefore the limb-darkening effects will be larger and more model dependent than is the case for data from the CHARA Array.

Our data for HD 222404 are on the second lobe of the visibility curve (Figure 3), and the second lobe is where second-order effects such as limb darkening start to have more of an influence than on the first lobe. In order to check that we are fitting only the angular diameter to these data and are not making unfair assumptions about the limb-darkening coefficient, we determined the diameter after changing μ_λ by 50%, which is well past the regime for stars of HD 222404’s general T_{eff} and $\log g$. The resulting change in diameters was $\sim 0.6\%$, indicating a low dependence on the μ_λ used.

Many of the stars in the sample are published in the literature as variable stars or as components in a binary star system. Table 5

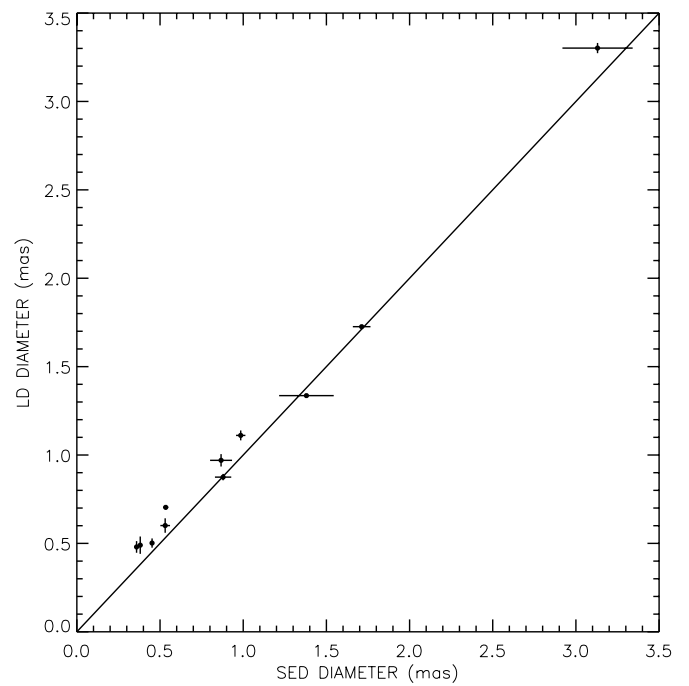


Figure 4. A comparison of estimated SED diameters and measured LD diameters with their corresponding errors. The solid line indicates a 1:1 ratio for the diameters. The LD diameter errors are consistently low, ranging between 0.01 to 0.05 mas, while the SED diameter errors show a wider spread, from 0.01 to 0.16 mas, and are dependent on how well the stellar model’s fluxes match the measured values. In the case of HD 221345 and HD 222404, which are the two points showing the largest SED errors, the model fluxes do not correspond as well to the measured fluxes.

lists the stars, the pertinent references, and why their variability or binarity do not affect our measurements here. For the variable stars, no reliable periods or types are listed in the literature, and if those stars are variable, it is on a level not likely to have a significant impact on our measurements. As for the binary star systems, the companions are too far away from the primary star and well out of the field of view (FOV) of the CHARA Array and/or the magnitude difference is too great for the Array to detect the secondary star.

The range of binary separations available to the CHARA Array, taking all the baselines into account, is approximately 10 mas to 1.0 arcsecond, while the maximum FOV of the baseline used for our observations is ~ 230 mas. The lower limit of binary detection using the CHARA Array is 2.5 magnitudes in the K band, and this value depends on the absolute brightness of the two stars and could therefore be higher for some systems. It is possible that the exoplanet parent stars may also host low-mass stellar companions not detected by the Array, though it is more likely they would have been detected by the radial velocity studies. We cannot detect the exoplanets themselves using the Array, due to the large magnitude difference between star and planet.

4. EFFECTIVE TEMPERATURES

Once θ_{LD} is measured, the effective temperature can be calculated using the relation

$$F_{\text{BOL}} = \frac{1}{4}\theta_{\text{LD}}^2\sigma T_{\text{eff}}^4, \tag{2}$$

where F_{BOL} is the bolometric flux and σ is the Stefan–Boltzmann constant. F_{BOL} was determined by applying the bolometric

Table 4
Exoplanet Host Star Angular Diameters and Radii

HD	Spectral Type	μ_λ	π (mas)	θ_{SED} (mas)	θ_{UD} (mas)	θ_{LD} (mas)	σ_{LD} (%)	R_{L} (R_\odot)	σ_{R} (%)
16141	G5 IV	0.27	25.67 ± 0.66	0.381 ± 0.012 ^b	0.480 ± 0.048	0.490 ± 0.049	10	2.05 ± 0.21	10
17092	K0 III	0.33	(183 ± 18 pc) ^a	0.531 ± 0.029 ^b	0.586 ± 0.039	0.601 ± 0.041	7	11.8 ± 1.4	12
45410	K0 III-IV	0.31	17.92 ± 0.47	0.867 ± 0.066	0.946 ± 0.034	0.970 ± 0.035	4	5.82 ± 0.26	4
154345	G8 V	0.28	53.80 ± 0.32	0.452 ± 0.008 ^b	0.490 ± 0.026	0.502 ± 0.026	5	1.00 ± 0.05	5
185269	G0 IV	0.25	19.89 ± 0.56	0.359 ± 0.012 ^b	0.471 ± 0.032	0.480 ± 0.033	7	2.59 ± 0.19	7
188310	G9 III	0.32	17.77 ± 0.29	1.712 ± 0.053	1.671 ± 0.008	1.726 ± 0.008	0.4	10.45 ± 0.18	2
199665	G6 III	0.31	13.28 ± 0.31	0.985 ± 0.028	1.083 ± 0.027	1.111 ± 0.028	3	9.00 ± 0.31	3
210702	K1 III	0.31	18.20 ± 0.39	0.879 ± 0.049 ^b	0.854 ± 0.017	0.875 ± 0.018	2	5.17 ± 0.15	3
217107	G8 IV	0.28	50.36 ± 0.38	0.534 ± 0.016 ^b	0.688 ± 0.013	0.704 ± 0.013	2	1.50 ± 0.03	2
221345	G8 III	0.32	12.63 ± 0.27	1.380 ± 0.164	1.297 ± 0.008	1.336 ± 0.009	1	11.38 ± 0.26	2
222404	K1 IV	0.32	70.91 ± 0.40	3.130 ± 0.211	3.331 ± 0.022	3.302 ± 0.029	1	5.01 ± 0.05	1

Notes.

^a HD 17092 had no parallax measurements available so we used the distance estimate from Gontcharov (2008). All spectral classes are from the *SIMBAD Astronomical Database*; μ_λ values are from Claret et al. (1995); π values are from van Leeuwen (2007).

^b θ_{SED} from van Belle & von Braun (2009); otherwise SEDs were completed using photometry from the following sources: HD 45410: *UBV* from Johnson et al. (1966), *RI* from Monet et al. (2003); HD 188310: *UBVRI* from Morel & Magnenat (1978); HD 199665: *BV* from Perryman & ESA (1997), *RI* from Monet et al. (2003); HD 221345: *UBV* from Johnson et al. (1966), *RI* from Monet et al. (2003); and HD 222404: *UBVRI* from Morel & Magnenat (1978). All *JHK* values from Cutri et al. (2003).

Table 5
Binary and Variable Stars in the Sample

Target HD	Type	Reference	Notes
16141	binary	Mugrauer et al. (2005)	$\rho = 6$ arcsec; outside Array's FOV ^a
45410	binary	Mason et al. (2001)	$\rho = 190$ arcsec; outside Array's FOV ^a
154345	variable	Samus et al. (2009)	no variability period or type listed
185269	binary	Strassmeier et al. (1989)	listed as binary but no orbital info given; no other indication in literature of binarity
188310	binary	Mason et al. (2001)	$\rho = 0.1$ arcsec, $\Delta m_V = 4.7$; outside range of Array
199665	binary	Mason et al. (2001)	$\rho = 130 - 200$ arcsec; outside Array's FOV ^a
217107	binary	Mason et al. (2001)	$\rho = 0.3 - 0.5$ arcsec; outside Array's FOV ^a
221345	variable	Hoffleit & Jaschek (1982)	no variability detected in Percy (1993)
222404	binary	Torres (2007)	$\rho = 325$ mas, $\Delta m_K = 6.4$; outside range of Array

Notes. ρ = binary separation, Δm = magnitude difference.

^a The field of view depends largely on the baseline used in the observations, so while some of the secondary companions would affect the data on shorter baselines, they will not be visible in the measurements on the baseline used here.

Table 6
Stellar Effective Temperatures and Luminosities

Star HD	A_V (mag)	BC	F_{BOL} (10^{-8} erg s ⁻¹ cm ⁻²)	Calculated T_{eff} (K)	$\sigma_{T_{\text{eff}}}$ (%)	Range of T_{eff} from other sources (K)	$\log(L)$ (L_\odot)
16141	0.00 ^a	0.06 ± 0.04	4.9 ± 0.2	4982 ± 254	5	4900-5888	2.3 ± 0.1
17092	0.80 ^a	0.50 ± 0.05	6.2 ± 0.4	4765 ± 182	4	4750	65.0 ± 3.1
45410	0.03 ^b	0.29 ± 0.03	15.2 ± 0.5	4689 ± 92	2	4750-4898	14.8 ± 0.4
154345	0.20 ^a	0.40 ± 0.04	8.6 ± 0.4	5664 ± 158	3	5436-5570	0.9 ± 0.0
185269	0.13 ^a	0.01 ± 0.03	6.0 ± 0.2	5283 ± 186	4	5850-6166	4.7 ± 0.1
188310	0.10 ^b	0.35 ± 0.02	50.2 ± 1.0	4742 ± 26	1	4635-4786	49.7 ± 0.9
199665	0.00 ^b	0.28 ± 0.04	26.8 ± 1.1	5054 ± 81	2	4750-5012	47.6 ± 1.8
210702	0.10 ^a	0.32 ± 0.03	14.2 ± 0.4	4859 ± 62	1	4600-4898	13.4 ± 0.4
217107	0.10 ^a	0.09 ± 0.03	9.5 ± 0.3	4895 ± 57	1	4900-5704	1.2 ± 0.0
221345	0.13 ^b	0.36 ± 0.03	32.3 ± 1.0	4826 ± 40	1	4582-4900	63.3 ± 1.8
222404	0.01 ^b	0.36 ± 0.00	184.0 ± 0.5	4744 ± 21	0.4	4566-4916	11.4 ± 0.0

Notes. All BC values from Allende Prieto & Lambert (1999) except for HD 17092 and HD 154345, which are from Cox (2000) with an assigned error of 10%. The range of T_{eff} values are from the *VizieR database of astronomical catalogs* (Ochsenbein et al. 2000).

^a van Belle & von Braun (2009).

^b Famaey et al. (2005).

corrections (BC) for each star after taking interstellar absorption (A_V) into account. Table 6 lists the A_V and BC used, and

the resulting F_{BOL} and T_{eff} . As a comparison, a range of T_{eff} from other sources is also listed in Table 6. Five stars have

Table 7
PARAM Model Results

Target HD	V mag	Average [Fe/H]	R_{model} (R_{\odot})	Mass (M_{\odot})	Age (Gyr)
16141	6.83	0.11 ± 0.07	2.3 ± 0.1	1.1 ± 0.0	7.2 ± 1.1
17092	7.82	0.00 ± 0.05	7.8 ± 0.4	1.5 ± 0.2	2.6 ± 0.9
45410	5.87	0.17 ± 0.05	6.1 ± 0.3	1.3 ± 0.1	4.0 ± 1.3
185269	6.70	0.11 ± 0.05	2.6 ± 0.1	1.4 ± 0.0	3.4 ± 0.2
188310	4.70	-0.27 ± 0.10	10.0 ± 0.4	1.0 ± 0.2	7.1 ± 3.6
199665	5.48	-0.10 ± 0.12	8.0 ± 0.3	2.0 ± 0.1	1.1 ± 0.1
210702	5.95	0.00 ± 0.05	5.2 ± 0.2	1.4 ± 0.1	3.5 ± 1.1
221345	5.22	-0.32 ± 0.05	10.3 ± 0.3	1.1 ± 0.2	4.5 ± 1.9
222404	3.21	0.08 ± 0.11	5.0 ± 0.2	1.2 ± 0.1	5.4 ± 2.1

Notes. V magnitudes are from Mermilliod (1991) except for HD 17092, which is from Droege et al. (2006); Average [Fe/H] are from the literature; R_{model} , Mass, and Age are model outputs.

T_{eff} within their ranges of temperatures obtained using other means, five stars are slightly out of their ranges but are within measured errors, and only one star is significantly outside its range (HD 185269, by ~ 570 K). This could be due to incorrect spectral typing or assumptions about factors such as opacity and metallicity that are buried in the model used for each of the three references that list temperatures for this star.

Because the θ_{LD} is dependent on the μ_{λ} value selected, which in turn is dependent on $\log g$ and T_{eff} , we wanted to check the effect of the new temperature values on measured LD diameters. Using the newly calculated T_{eff} to find μ_{λ} , we found the average difference in μ_{λ} was $<6\%$ and the resulting θ_{LD} values differed on average of 0.3%, indicating this is a negligible effect.

5. STELLAR MODEL RESULTS

In order to estimate stellar ages, masses, and linear radii, we used the PARAM 1.0 model⁶ (da Silva et al. 2006), which is based on a set of theoretical isochrones from Girardi et al. (2000). The model uses each star’s metallicity, effective temperature, and V magnitude to estimate its age, mass, radius, $(B - V)_0$, and $\log g$ using the isochrones and a Bayesian estimating method, calculating the probability density function separately for each property in question. da Silva et al. are most confident in resulting $(B - V)_0$, $\log g$, radii, and angular diameter predictions while describing the age and mass estimates as “more uncertain.” We left the Bayesian priors (initial mass function and star formation rate in a given interval) on the default settings when running the model.

The model’s inputs were the star’s T_{eff} , [Fe/H], V magnitude, and parallax along with the corresponding error for each value. T_{eff} was calculated using Equation (3), the V magnitude was from Mermilliod (1991), the parallax was from van Leeuwen (2007), and the [Fe/H] value was averaged from all the sources available from Ochsenein et al. (2000) with its error represented by the standard deviation of all the measurements. When only one source of [Fe/H] was in the literature (the case for HD 45410 and HD 185269), an error of 0.05 was assigned. The same error was used when the star had no [Fe/H] listed and solar metallicity was assumed (the case for HD 17092, HD 154345, and HD 210702).

The resulting age, mass, and R_{model} are listed in Table 7 for all the stars except HD 154345 because it is a dwarf and the model is for evolving stars, and for HD 217107, whose metallicity is out of range of the model. Figure 5 plots the model’s radii versus

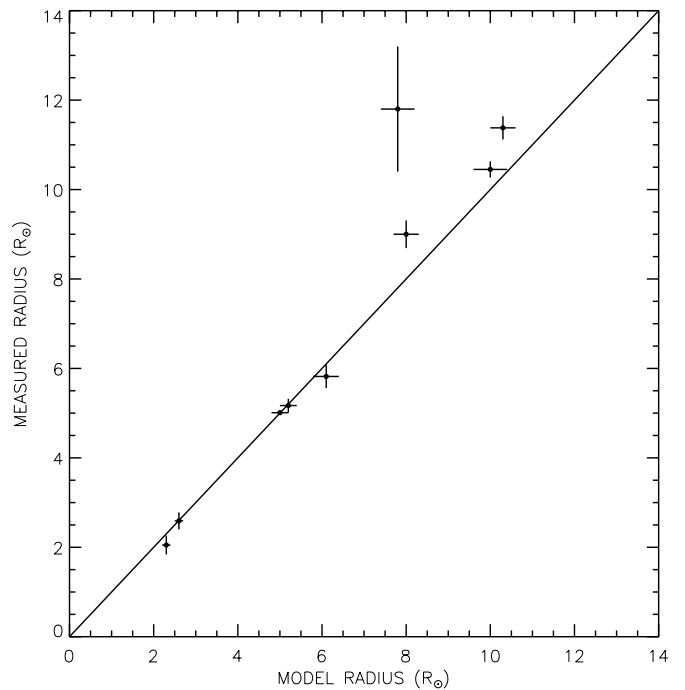


Figure 5. Comparison of model and measured radii with their corresponding errors. The solid line indicates a 1:1 ratio for the radii. The measured radii errors depend on uncertainties in the LD diameter and parallax measurements while the model radii errors depend on the model’s inputs, including effective temperature, metallicity, and parallax measurements. The errors in each input value contribute to the error budget of the model radius. The largest outlier is HD 17092, which had the least reliable distance measurement of the sample.

those measured interferometrically. The agreement between the two is excellent for the small to intermediate-sized stars, but the model appears to systematically underestimate the radii for the four largest stars. Figure 6 plots luminosity versus T_{eff} and represents the Hertzsprung–Russell (H–R) diagram. The zero-age main-sequence (ZAMS) line is shown as derived from Cox (2000) and the one dwarf in the sample (HD 154345) is the point nearly on the ZAMS while the other stars form the giant branch.

6. CONCLUSION

We measured the angular diameters of 11 exoplanet host stars for a sample almost entirely comprised of evolving and evolved stars. All LD diameters boasted errors of $\leq 10\%$, and 8 of the 11 had errors $\leq 5\%$. Linear radii were derived from θ_{LD} and the stars’ *Hipparcos* measurements, and we calculated effective

⁶ http://stev.oapd.inaf.it/cgi-bin/param_1.0

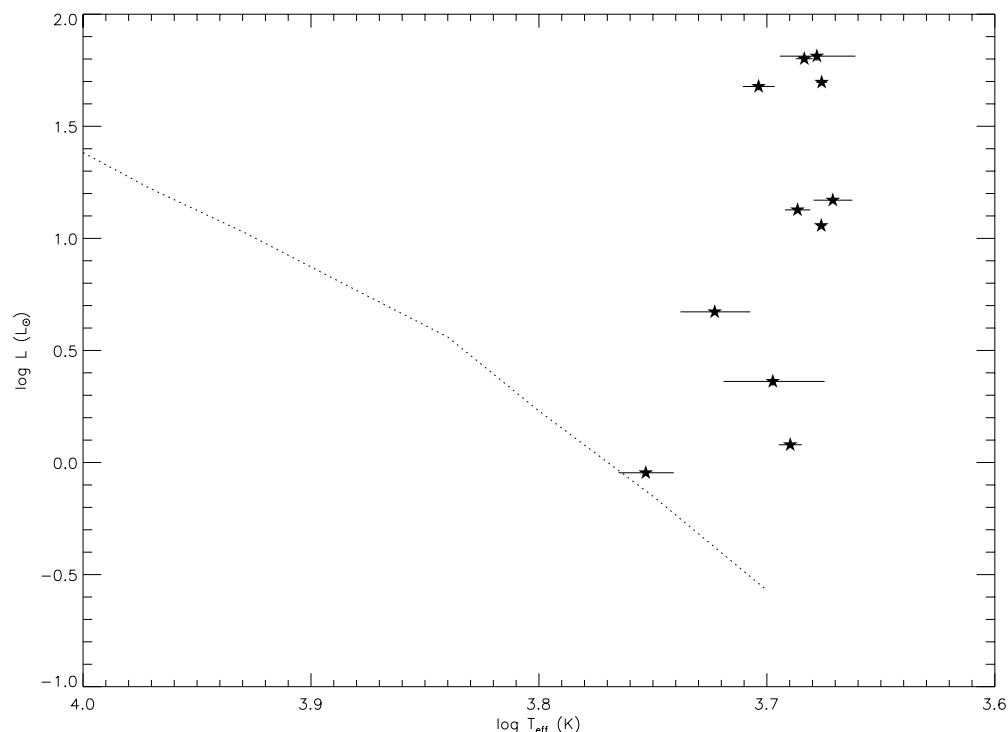


Figure 6. H–R diagram for the exoplanet host stars. The dotted line indicates the ZAMS derived from Cox (2000). The star closest to this line is HD 154345 and is the only dwarf in the sample. The remaining points represent the giant branch of the H–R diagram. The main sources of error in the luminosity values arise from uncertainties in bolometric corrections (the error bars are within the data points), while the effective temperature errors depend on uncertainties in the star’s parallax and LD diameter measurements.

temperatures using our θ_{LD} values. The subsequent errors on the T_{eff} were all $\leq 5\%$.

Using our new effective temperatures, [Fe/H] values from the literature, and the PARAM stellar model, we were able to estimate the radii, masses, and ages for the stars, and the model radii match the measured radii well for all the giants except the four largest stars in the sample. Previous interferometric measurements of other giant stars showed a similar effect, where high-luminosity stars have larger radii at a given effective temperature (Dyck et al. 1998). The four stars in question—HD 17092, HD 188310, HD 199665, and HD 221345—are by far the most luminous stars in the sample so it is not entirely unexpected that the models underestimate their radii. It would be to the model’s advantage if it could be modified to incorporate this effect.

By directly measuring exoplanet host stars’ angular diameters and calculating the physical radii and temperatures, we are able to better characterize the exoplanets’ environments. We now know that solar systems come in many different configurations (Butler et al. 2006), and interferometric measurements help to describe the all-important central stars. This in turn will help to constrain parameters such as the location and size of the habitable zone as well as putting limitations on the temperature profiles of the planets themselves.

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